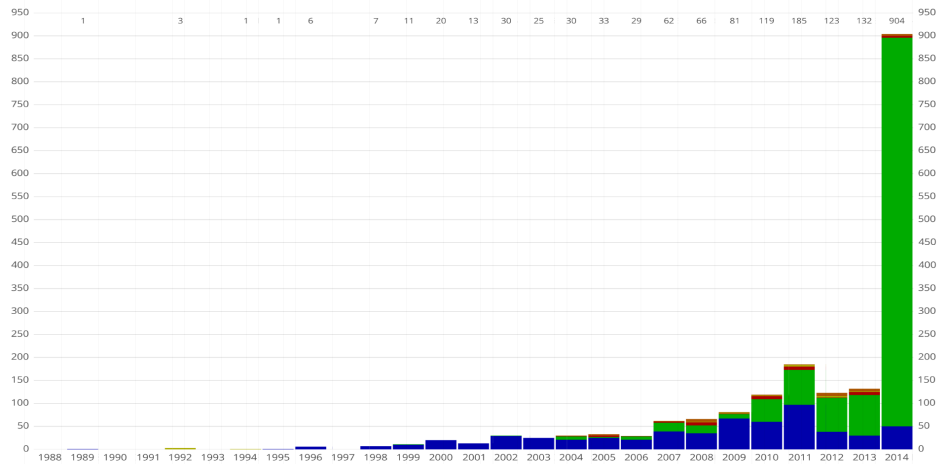
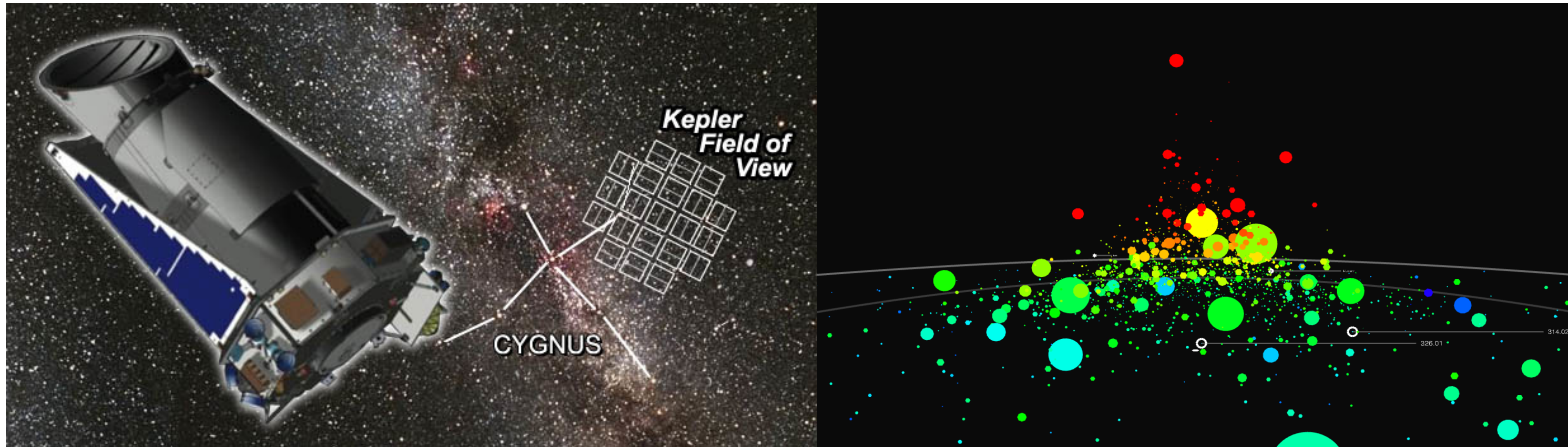


Least squares fitting from linear algebra enables identification of exoplanets

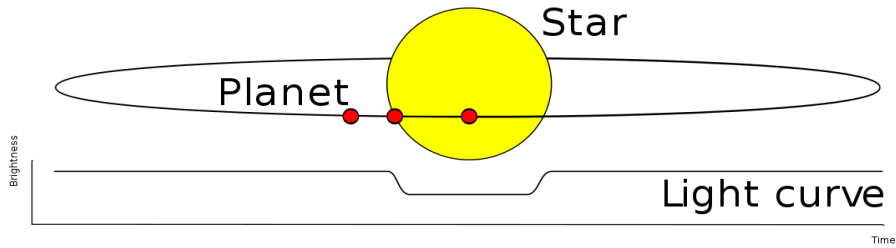
- Exoplanet search techniques
- Fitting transit curve

- Kepler is a NASA exoplanet search mission

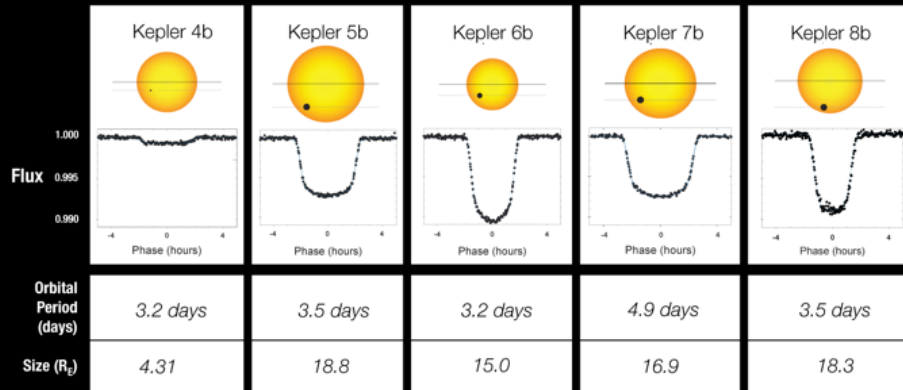


- Direct imaging
- Transit
- Radial velocity
- Microlensing
- Timing

All require application of linear algebra in order to identify signal

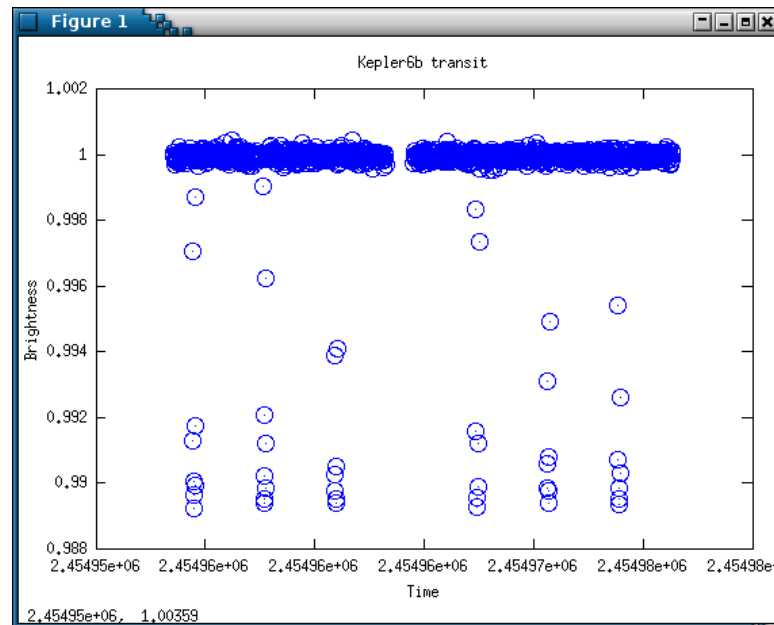


## Transit Light Curves

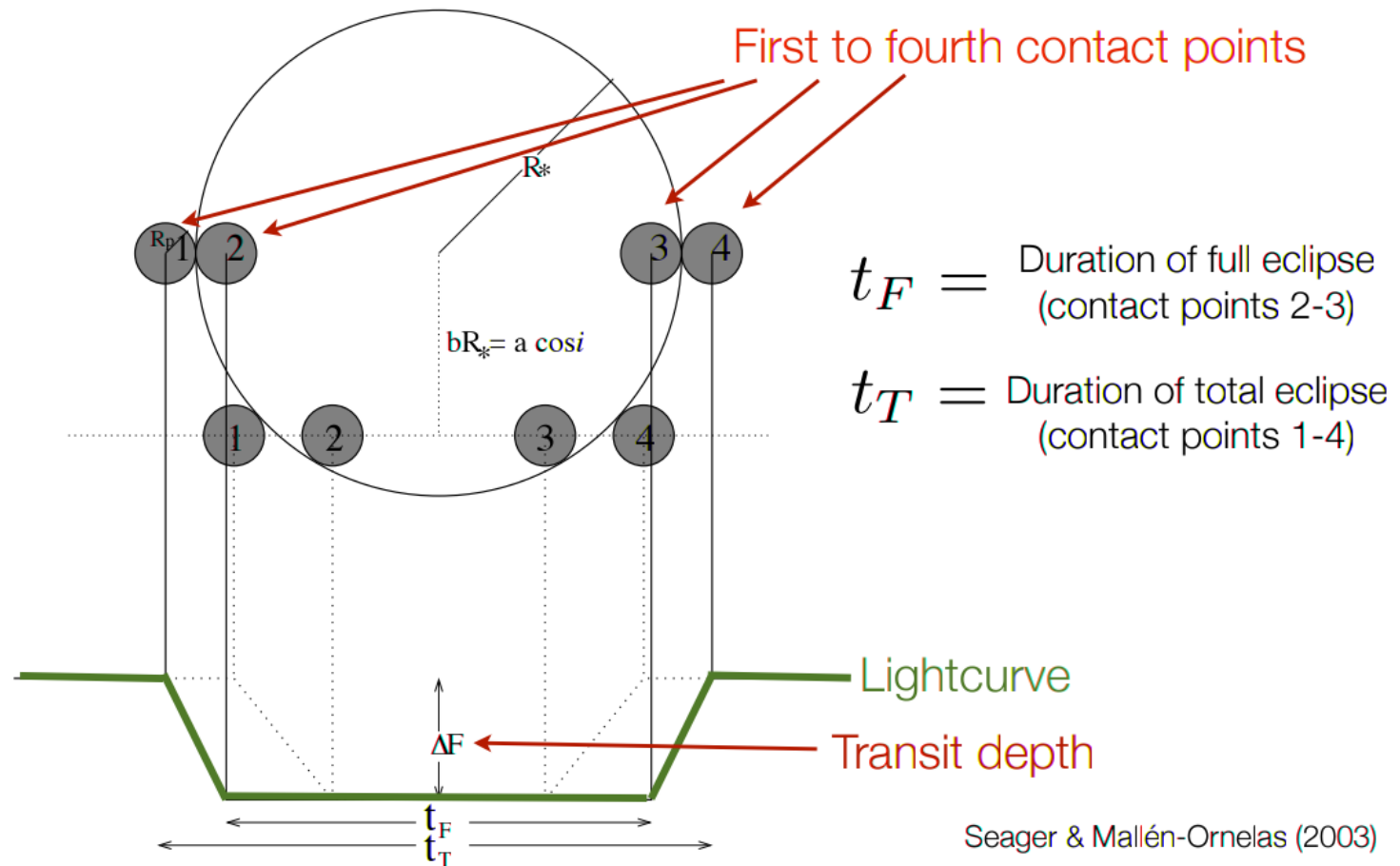


# Least squares fit

```
octave> cd /home/student/courses/MATH547/lessons/exoplanets;  
octave> data=textread('Kepler6b.data');  
octave> m=max(size(data))/2; t=data(1:2:m-1); b=data(2:2:m);  
octave> plot(t,b,'ob'); xlabel('Time'); ylabel('Brightness'); title('Kepler6b transit');  
octave> print -dpng kep6bdata.png;  
octave>
```



# Solving the star-planet system



# Solving the star-planet system

From Kepler's third law and radial velocity we have:

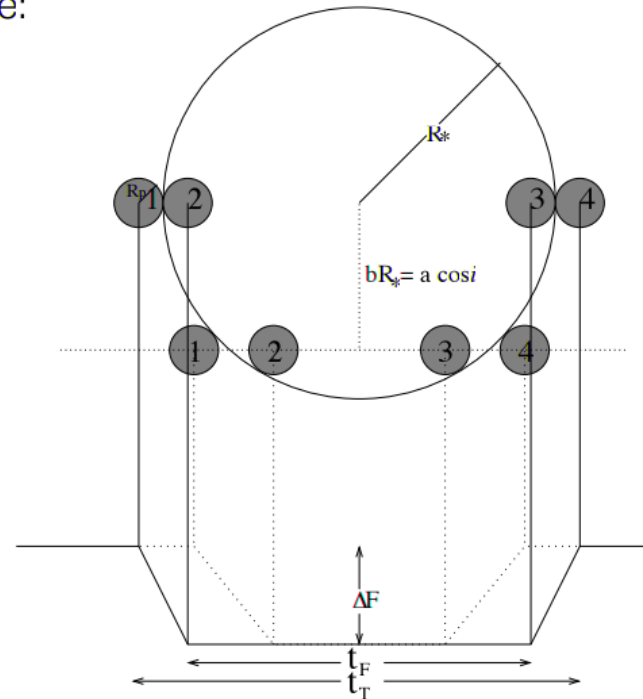
$$P^2 = \frac{4\pi^2 a^3}{G(M_* + M_p)}$$

From the transit depth measurement we have the planet/star ratio:

$$\frac{R_p}{R_*} = \sqrt{\Delta F}$$

We can now calculate the orbital inclination:

$$i = \cos^{-1} \left( b \frac{R_*}{a} \right)$$



Seager & Mallén-Ornelas (2003)

# Solving the star-planet system

We can now calculate the orbital inclination:

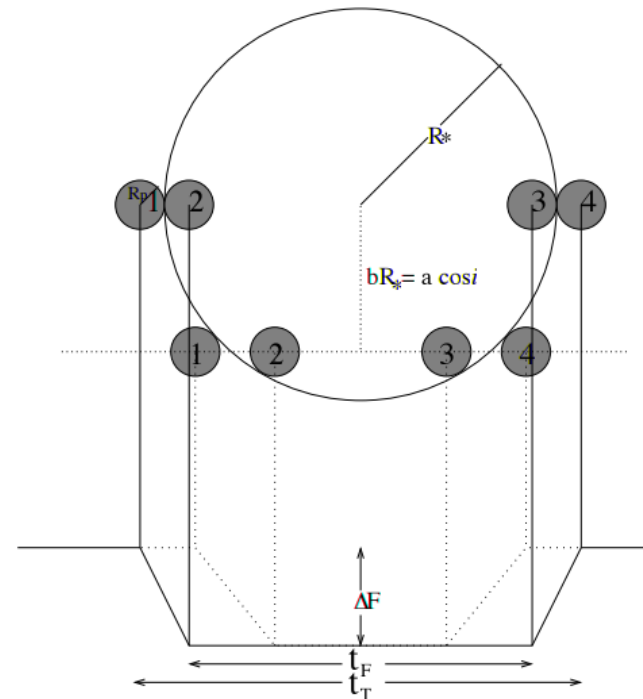
$$i = \cos^{-1} \left( b \frac{R_*}{a} \right)$$

What is  $b$ ?? - it's called the impact parameter:

$$b = \left[ \frac{\left(1 - \sqrt{\Delta F}\right)^2 - (t_F/t_T)^2 \left(1 + \sqrt{\Delta F}\right)^2}{1 - (t_F/t_T)^2} \right]^{1/2}$$

We can even calculate the stellar density:

$$\rho_* = \frac{32}{G\pi} P \frac{\Delta F^{3/4}}{(t_T^2 - t_F^2)^{3/2}}$$



Seager & Mallén-Ornelas (2003)



```
octave>
```